

Quick Guidances

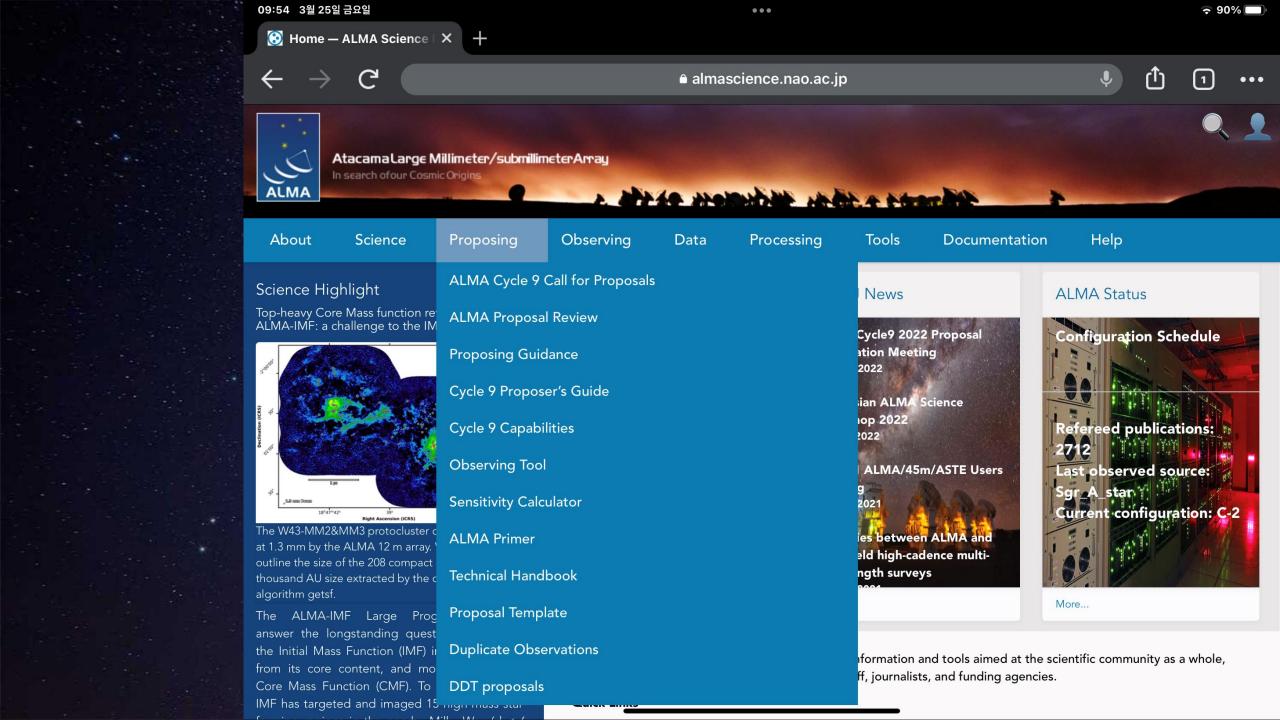
Before using OT

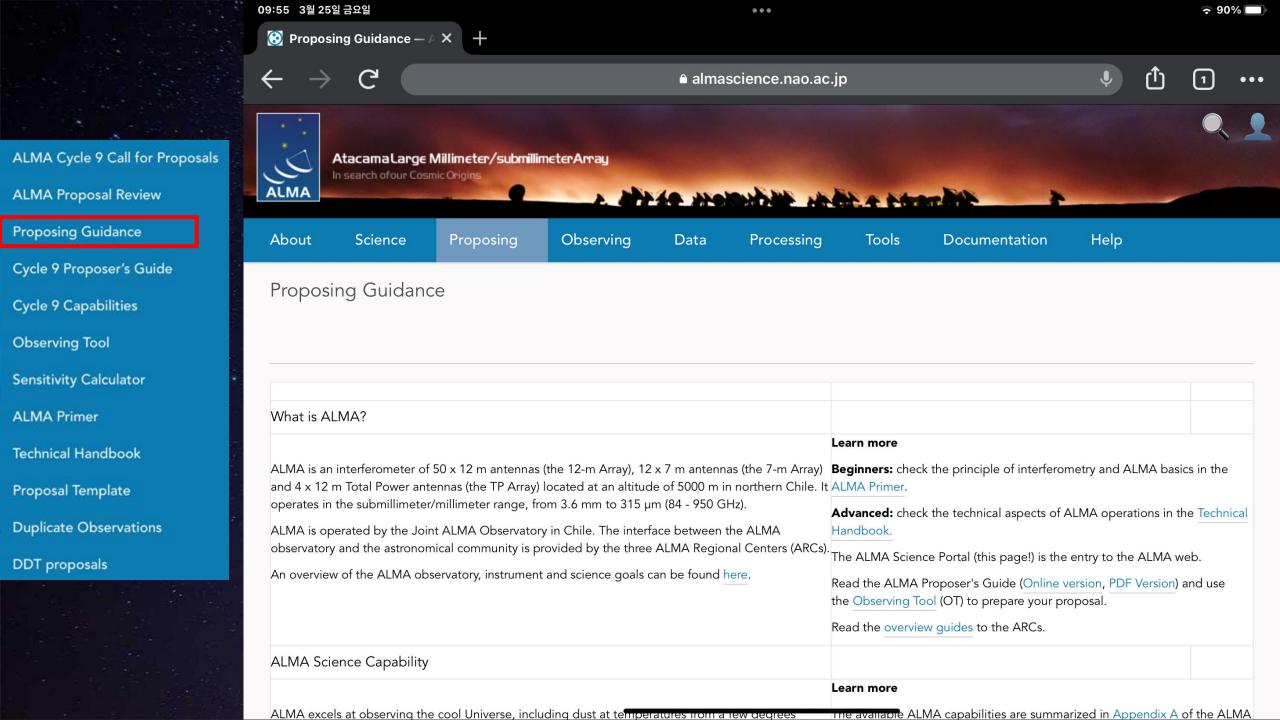
1 Important Input Parameters
Sclaes & Spectral Set up.

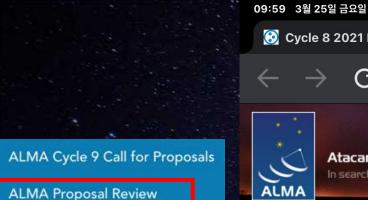
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# ALMA Proposal Review

The ALMA proposal review process is organized by the Proposal Handling Team (PHT) at the Joint ALMA Observatory (JAO). ALMA proposals are selected through competitive peer review through either the distributed peer review process or the ALMA Proposal Review Committee (APRC).

Proposals that request less than 50 h on the 12-m Array or less than 150 h on the 7-m Array in standalone mode are reviewed using the distributed peer review system, in which the proposal team designates one member of the proposal team to participate in the review process. The outcomes of the distributed peer review process are:

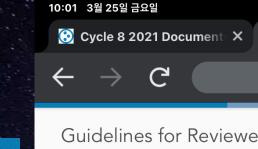
- A scientifically ranked list of proposals
- Individual comments for each proposal written by the reviewers that are sent to the Principal Investigators (PIs).

Large Programs, i.e., proposals that request more than 50 h on the 12-m Array or more than 150 h on the 7-m Array in standalone mode, are reviewed by the APRC, which is a panel composed of experts selected from the international astronomical community. To gain further expert advice, external Science Assessors will provide reviews on Large Programs, which will be considered by the APRC. The outcomes of the APRC review process are:

- A list of recommended Large Programs
- A consensus report for each Large Program that summarizes the strength and weaknesses of the proposal.

All proposals are reviewed in a dual anonymous fashion in which the proposers do not know the identity of the reviewers and the reviewers do not know the identity of the proposers. All proposals need to be prepared in accordance with the dual-anonymous guidelines.

Available documents for the proposal review process





Additional review criteria for Large Programs

Review criteria

Conflict criteria

Unconscious bias

**Dual-anonymous** 

Writing reviews to the Pls

**Guidelines for Mentors** 

Code of conduct and confidentiality





# ALMA Cycle 9 Call for Proposals

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# Guidelines for Reviewers

# Review criteria

Each proposal contains a cover sheet, a Scientific Justification, and a Technical Justification. Reviewers need to read each of these sections. Note in particular that the Technical Justification often contains a detailed justification of the requested sensitivity, angular resolution, and correlator setup that will be useful in evaluating the proposal.

**⊗** Guidelines for Reviewer∈ ×

Reviewers should assess the scientific merit of the proposals to the best of their ability using the following criteria:

The overall scientific merit of the proposed investigation and its potential contribution to the advancement of scientific knowledge.

- Does the proposal clearly indicate which important, outstanding questions will be addressed?
- Will the proposed observations have a high scientific impact on this particular field and address the specific science goals of the proposal? ALMA encourages reviewers to give full consideration to well-designed high-risk/high-impact proposals even if there is no guarantee of a positive outcome or definite detection.

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• Does the proposal clearly describe how the data will be analyzed in order to achieve the science goals?

The suitability of the observations to achieve the scientific goals.

- Is the choice of target (or targets) clearly described and well justified?
- Are the requested signal-to-noise ratio, angular resolution, largest angular scale, and spectral setup sufficient to achieve the science goals and well justified?
- Does the proposal justify why new observations are needed to achieve the science goals?

In general, the scientific merit should be assessed solely on the content of the proposal, according to the above criteria. Proposals may contain references to published papers (including preprints) as per standard practice in the scientific literature. Consultation of those references should not, however, be required for a general understanding of the proposal.

Technical feasibility.

The ALMA Observing Tool (OT) validates most technical aspects of the proposal; e.g., the OT verifies that the angular resolution can be achieved, verifies the correlator setup is feasible, and provides an accurate estimate of the integration time needed to achieve the requested sensitivity. Reviewers should assume that the OT technical validation of the proposal is correct. Reviewers should not downgrade a proposal merely because of the amount of time requested by the proposal. However, the reviewers can and should consider if the requested signal-to-noise ratio, angular resolution, largest angular scale, and spectral setup as requested by the PI are sufficient to achieve the scientific goals of the proposal and are well justified.



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# Observing Tool

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The ALMA Observing Tool (OT) is a Java desktop application used for the preparation and submission of ALMA Phase 1 proposals and, for those which are accepted, Phase 2 materials (Scheduling Blocks). It is also used for preparing and submitting Director's Discretionary Time (DDT) proposals and Supplemental Call (ACA stand-alone) proposals. The current Cycle 9 release of the OT is configured for the present capabilities of ALMA as described in the Cycle 9 Call For Proposals. Note that in order to submit proposals you will have to register with the ALMA Science Portal beforehand.

## Download & Installation

Documentation

The OT should run on all common operating systems and depends on a version of Java being available. In previous releases of the OT it was the responsibility of the user to ensure that a suitable version of Java was installed, but the Cycle 9 version of the OT will come with its own version of Java 11 and thus the users need no longer worry about their local Java installation. Unfortunately, as Java 11 does not include Web Start, this version of the OT is no longer available. The Cycle 9 OT can be installed in two different ways, either with a modern installer or manually with a tarball distribution.

It is recommended that the OT be installed using the ALMA OT Installer. This uses a modern graphical interface to report the progress of the installation and allows the user to change various settings from their defaults, including the amount of memory the OT may use. The installation will produce an executable file that can be used to start the OT. With the loss of Web Start, automatic updates of the tool are no longer possible, but the OT will detect if an update is available at start-up and inform the user. If problems are encountered with the installer, then the tarball must be used.

The tarball version must be installed manually and the instructions for doing this have not changed.

Installer Tarball 09:55 3월 25일 금요일 ● almascience.nao.ac.jp

Data

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**Processing** 



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Installer Page

Atacama Large Millimeter/submillimeterArray

Proposing

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In search of our Cosmic Origins

Science

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Mac OS Installer

Tools

Help

Documentation

- Linux Installer
- Windows Installer

Click on one of the links next to the OT Logo to download the Cycle-9 OT Installer for your particular operating system. The Installer is an executable file which can be started by double-clicking in a file-manager window or started from a shell's command line. Once started, it will take you through a number of screens which, for example, allow you to change the default amount of memory available to the OT. In most cases you can just accept all the defaults using the 'Next' button and click 'Install' when you are happy.

After the Installer has finished, an executable file ('ALMA-OT.sh' on Linux and 'ALMA-OT.app' on Macs) should be found inside a directory named 'ALMAOT-C9-2022'. This can be run from the command line or by double-clicking in a file manager if this is configured in this way. We recommend that the name of this directory not be changed so that multiple versions of the OT (for use in different cycles) can be maintained on your computer. On Macs, a shortcut will be created on your Desktop with the name 'ALMAOT-C9-2022' - the OS will probably ask to control your Finder for this to happen.

### Additional Information

- The Mac download is a zip archive which must first be opened in order to extract the installer. This will often be done automatically for you or a suitable program will be suggested ('Archive Utility').
- On Linux, typing 'sh almaot-C9-2022.bin' is the recommended way of starting the installer it should not be necessary to make it executable.
- There may be various issues related to security when running the Installer. Mac users may need to give permission to run the tool by opening the 'Security & Privacy' menu of 'System Preferences' and this menu should also be set to allow the use of apps from 'identified developers'. Alternatively, running the installer by right-clicking and choosing 'Open' (maybe twice) might work. On Windows, we are aware of 'Defender SmartScreen' this can be bypassed by clicking on 'More Info'.
- It also appears that the installer will not work on older versions of macOS. So far, we only know that this is the case for 10.10 Yosemite. Users of this OS will have to use the tarball.
- In contrast to the previous 'automated' OT installation (Web Start), the OT will no longer update itself automatically if an update is released. However, the OT will inform

YOU (and your Co-ls) shoud be registered in ALMA site.

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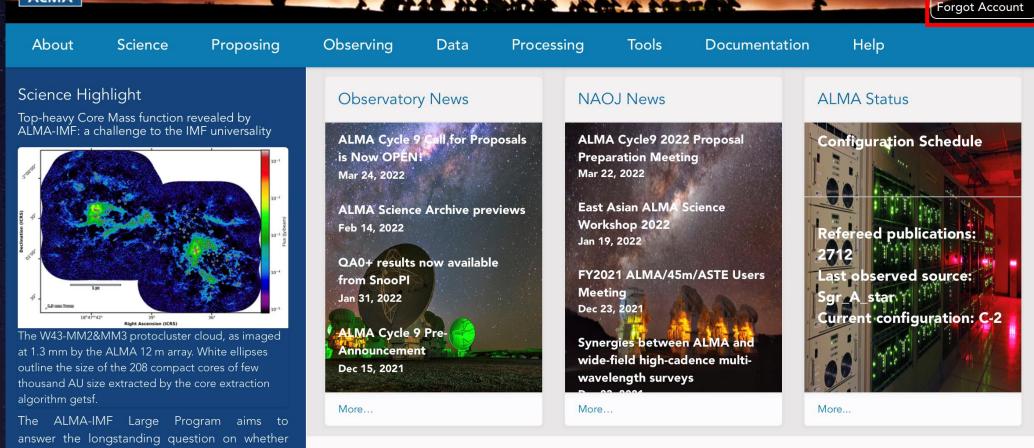
the Initial Mass Function (IMF) inherits its shape

from its core content, and more precisely the

Core Mass Function (CMF). To do that, ALMA-IMF has targeted and imaged 15 high-mass star-



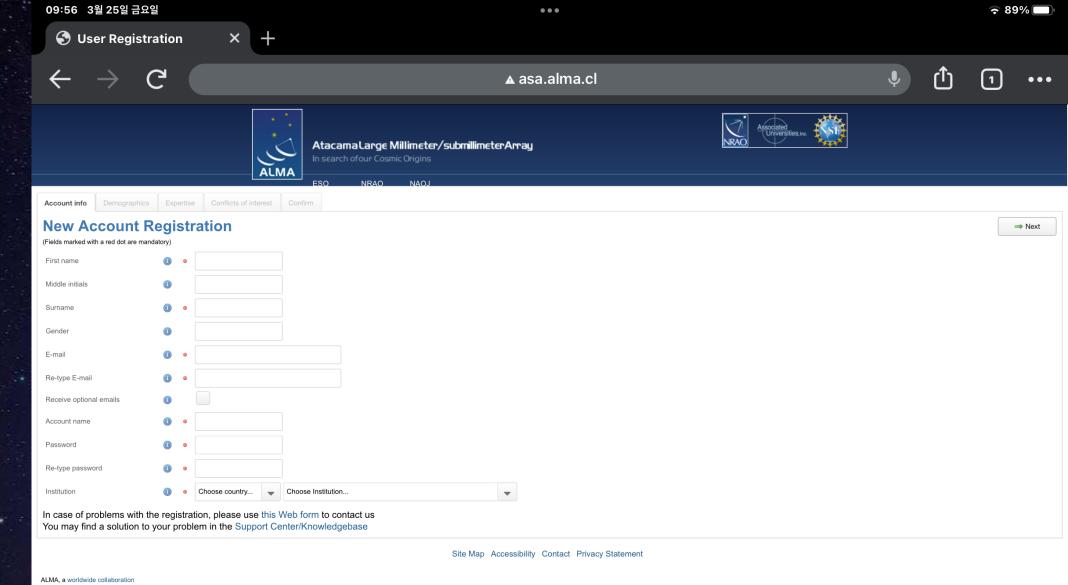
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The ALMA Science Portal is a one-stop source for information and tools aimed at the scientific community as a whole, including proposers, archive researchers, ALMA staff, journalists, and funding agencies.

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**Quick Links** 



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# Important Input Parameters

# Important parameters I

- Scales
  - Angular Resolution (beam size) ~ depend on the longest baseline
  - Maximum Recoverable Scale (MRS)
    - depends on the shortest baseline ( ~ 10 x beam size)
    - When the scale is loger than MRS, the emission is resolve out
    - Largest Angular Structure (LAS) should be shorther than MRS.
    - LAS > MRS → multiple configuration or ACA and TP are added.
  - Field of View (FOV)
    - FWHM of the 12m telescope primary beam
    - Area of target is larger than 1/3 FOV, mosaic is needed.

Table 2: Planned 12-m Array Configuration Schedule for Cycle 9

Start date	Configuration	Longest baseline	LST for best observing conditions		
2022 October 1	C-3	0.50 km	~ 22—10 h		
2022 October 20	C-2	0.31 km	~ 23—11 h		
2022 November 10	C-1	0.16 km	~ 1—13 h		
2022 November 30	C-2	0.31 km	~ 2—14 h		
2022 December 20	C-3	0.50 km	~ 4—15 h		
2023 January 10	C-4	0.78 km	~ 5—17 h		
2023 February 1	No observations due to maintenance				
2023 March 1	C-4	0.78 km	~ 8—21 h		
2023 March 20	C-5	1.4 km	~ 9—23 h		
2023 April 20	C-6	2.5 km	~ 11—1 h		
2023 May 20	C-7	3.6 km	~ 13—3 h		
2023 June 20	C-8	8.5 km	~ 15—5 h		
2023 July 11	C-9	13.9 km	~16—6 h		
2023 July 30	C-10	16.2 km	~17—7 h		
2023 August 20	C-9	13.9 km	~19—8 h		
2023 September 10	C-8	8.5 km	~20—9 h		

Table A-1: Angular Resolutions (AR) and Maximum Recoverable Scales (MRS) for the Cycle 9 configurations

Config	Lmax		Band 3	Band 4	Band 5	Band 6	Band 7	Band 8	Band 9	Band 10
	Lmin		100 GHz	150 GHz	185 GHz	230 GHz	345 GHz	460 GHz	650 GHz	870 GHz
7-m	45 m	AR	12.5"	8.35"	6.77"	5.45"	3.63"	2.72"	1.93"	1.44"
	9 m	MRS	66.7"	44.5"	36.1"	29.0"	19.3"	14.5"	10.3"	7.67"
C-1	161 m	AR	3.38"	2.25"	1.83"	1.47"	0.98"	0.74"	0.52"	0.39"
	15 m	MRS	28.5"	19.0"	15.4"	12.4"	8.25"	6.19"	4.38"	3.27"
C-2	314 m	AR	2.30"	1.53"	1.24"	1.00"	0.67"	0.50"	0.35"	0.26"
	15 m	MRS	22.6"	15.0"	12.2"	9.81"	6.54"	4.90"	3.47"	2.59"
C-3	500 m	AR	1.42"	0.94"	0.77"	0.62"	0.41"	0.31"	0.22"	0.16"
	15 m	MRS	16.2"	10.8"	8.73"	7.02"	4.68"	3.51"	2.48"	1.86"
C-4	784 m	AR	0.92"	0.61"	0.50"	0.40"	0.27"	0.20"	0.14"	0.11"
	15 m	MRS	11.2"	7.50"	6.08"	4.89"	3.26"	2.44"	1.73"	1.29"
C-5	1.4 km	AR	0.55"	0.36"	0.30"	0.24"	0.16"	0.12"	0.084"	0.063"
	15 m	MRS	6.70"	4.47"	3.62"	2.91"	1.94"	1.46"	1.03"	0.77"
C-6	2.5 km	AR	0.31"	0.20"	0.17"	0.13"	0.089"	0.067"	0.047"	0.035"
	15 m	MRS	4.11"	2.74"	2.22"	1.78"	1.19"	0.89"	0.63"	0.47"
C-7	3.6 km	AR	0.21"	0.14"	0.11"	0.092"	0.061"	0.046"	0.033"	0.024"
	64 m	MRS	2.58"	1.72"	1.40"	1.12"	0.75"	0.56"	0.40"	0.30"
C-8	8.5 km	AR	0.096"	0.064"	0.052"	0.042"	0.028"	0.021"	0.015"	0.011"
	110 m	MRS	1.42"	0.95"	0.77"	0.62"	0.41"	0.31"	0.22"	0.16"
C-9	13.9 km	AR	0.057"	0.038"	0.031"	0.025"	0.017"	0.012"	0.0088"	N/A
	368 m	MRS	0.81"	0.54"	0.44"	0.35"	0.24"	0.18"	0.13"	
C-10	16.2 km	AR	0.042"	0.028"	0.023"	0.018"	0.012"	0.0091"	N/A	N/A
	244 m	MRS	0.50"	0.33"	0.27"	0.22"	0.14"	0.11"		

Table A-2: Allowed Array Combinations and Time Multipliers

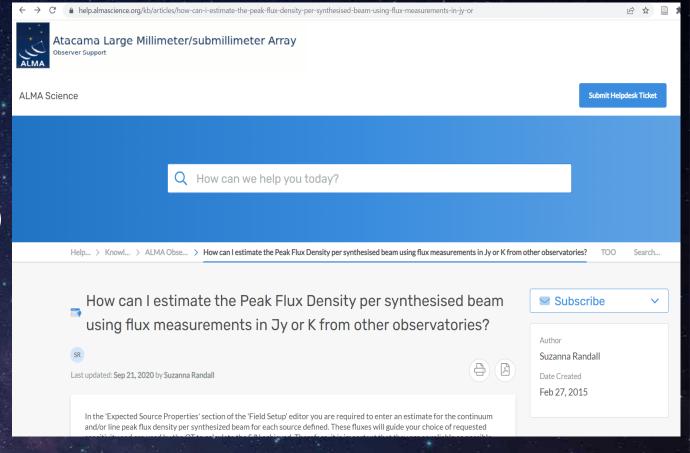
Most Extended configuration	Allowed Compact configuration pairings	Extended 12-m Array Multiplier	Multiplier if compact 12-m Array needed	Multiplier if 7-m Array needed	Multiplier if TP Array needed and allowed
7-m Array	TP			1	1.7
C-1	7-m Array & TP	1		7.0	11.9
C-2	7-m Array & TP	1		4.7	7.9
C-3	7-m Array & TP	1		2.4	4.1
C-4	C-1 & 7-m Array & TP	1	0.34	2.4	4.0
C-5	C-2 & 7-m Array & TP	1	0.26	1.2	2.1
C-6	C-3 & 7-m Array & TP	1	0.25	0.6	1.0
C-7	C-4	1	0.23		
C-8	C-5	1	0.22		
C-9	C-6	1	0.21		
C-10	-	1			

## Notes for Table A-2:

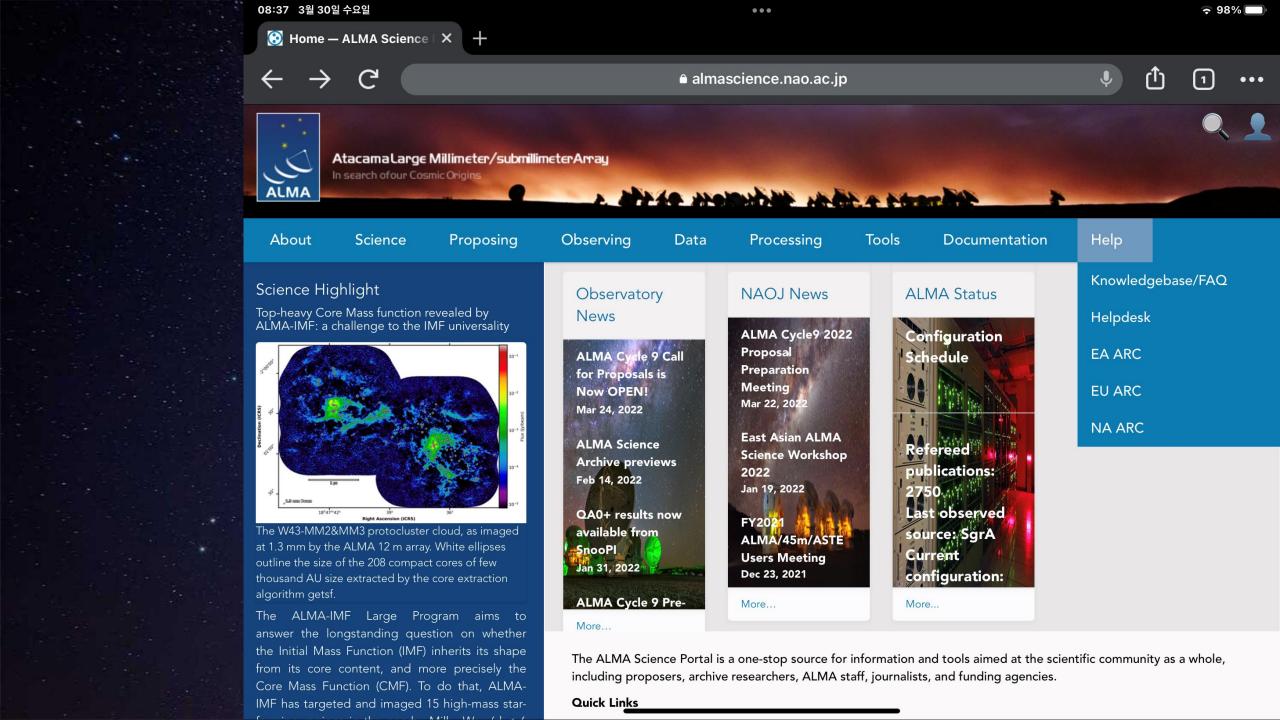
- See Chapter 7 of the <u>Technical Handbook</u> for relevant equations and detailed considerations.
   If the array configuration that meets the AR request according to Table A-1 has a MRS that is smaller than the LAS request, the OT checks if adding more compact array configurations, following the restrictions of this Table, fulfills the LAS request. If so, the final setup consists of the selected combination of arrays. Otherwise, the OT returns a validation error.

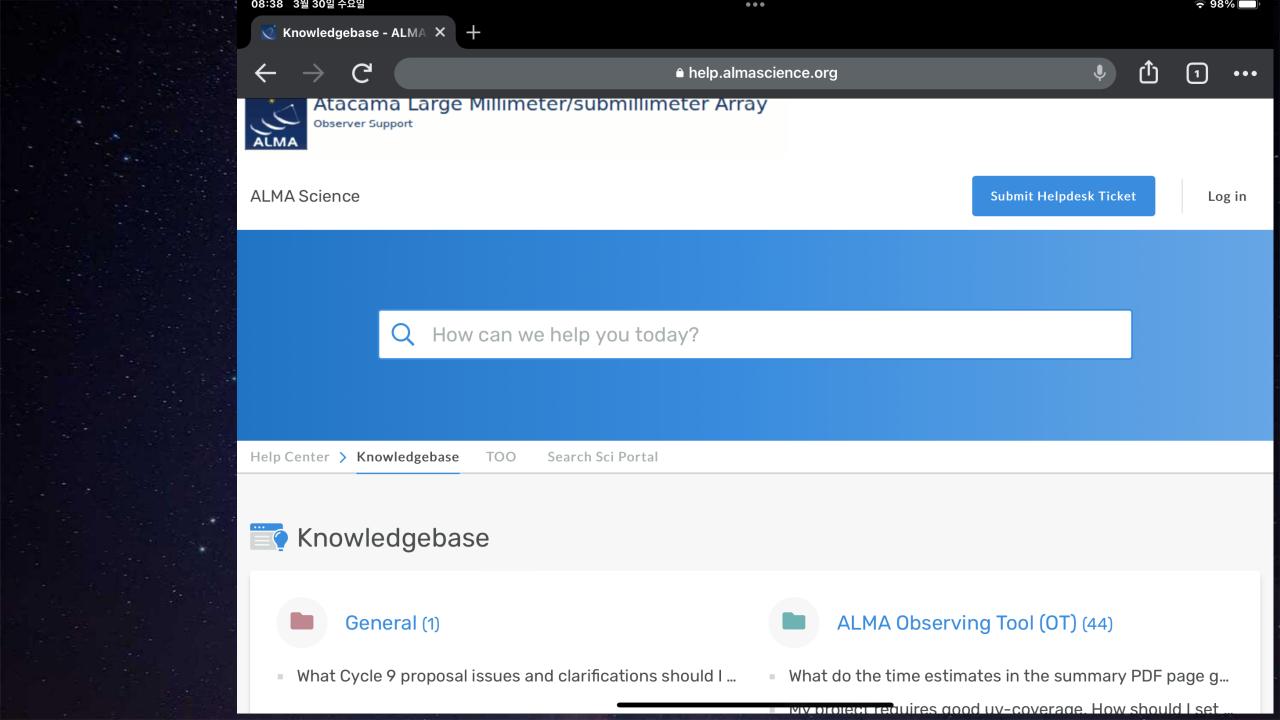
# Important parameters II

- Expected Source properties
  - Position, souce velocity
  - Peak Flux Density per beam
    - SNR > 3
  - Polarization
    - linear > 0.1% ( < 0.3 FOV)</li>
    - circular > 1.8 % (<0.1 FOV)
  - Line width
    - > 3 x spectral resolution



https://help.almascience.org/kb/articles/how-can-i-estimate-the-peak-flux-density-per-synthesised-beam-using-flux-measurements-in-jy-or





# Important parameters III

- Spectral Setup
  - LSB and/or USB
  - 4 basebands (with 2GHz max. width)
  - 2 or 4 basebands in the one sideband

LSB/USB (<4GHz, ≤4 basebands)
Baseband (<2GHz, ≤4spws)
Spectral window(spw)

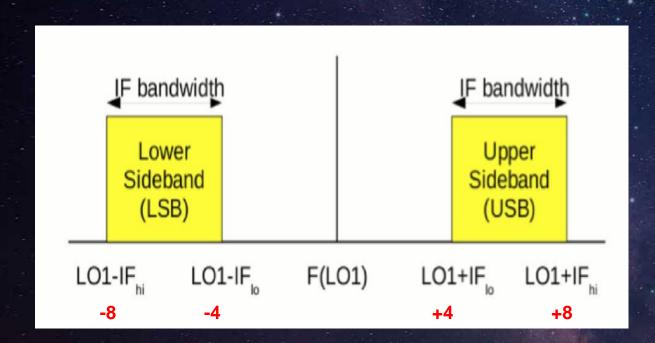


Table A-3: Properties of ALMA Cycle 9 Receiver Bands

Band	Frequency range <sup>1</sup> (GHz)	Wavelength range (mm)	IF range (GHz)	Туре
3	84 - 116	3.6 - 2.6	4 – 8	2SB
4	125 – 163	2.4 – 1.8	4 – 8	2SB
5	158 – 211	1.9 – 1.4	4 – 8	2SB
6	211 – 275	1.4 – 1.1	4.5 – 10	2SB
7	275 – 373	1.1 – 0.8	4 – 8	2SB
8	385 – 500	0.78 - 0.60	4 – 8	2SB
9	602 – 720	0.50 - 0.42	4 – 12	DSB
10	787 – 950	0.38 - 0.32	4 – 12	DSB

### Notes for Table A-3:

 These are the nominal frequency ranges for continuum observations. Observations of spectral lines that are within about 0.2 GHz of a band edge are not possible (at present) in Frequency Division Mode (FDM, see Section A.6.1) because of the responses of the spectral edge filters implemented in the correlator. IF is the intermediate frequency. Spws in a baseband

- one faction 1
- two fraction ½
- four faction 1/4
- one fraction ½ + two fraction ¼

Spectral windows (SPW) should have the same resolution.

Table A-4: Properties of ALMA Cycle 9 Correlator Modes, dual-polarization operation 1,2

Bandwidth (MHz)	Channel spacing <sup>(3)</sup> (MHz)	Spectral resolution (MHz)	Number of channels	Correlator mode <sup>(4)</sup>
1875	15.6	31.2	120	TDM
1875	0.488	0.976	3840	FDM
938	0.244	0.488	3840	FDM
469	0.122	0.244	3840	FDM
234	0.061	0.122	3840	FDM
117	0.0305	0.061	3840	FDM
58.6	0.0153	0.0305	3840	FDM

### Notes for Table A-4:

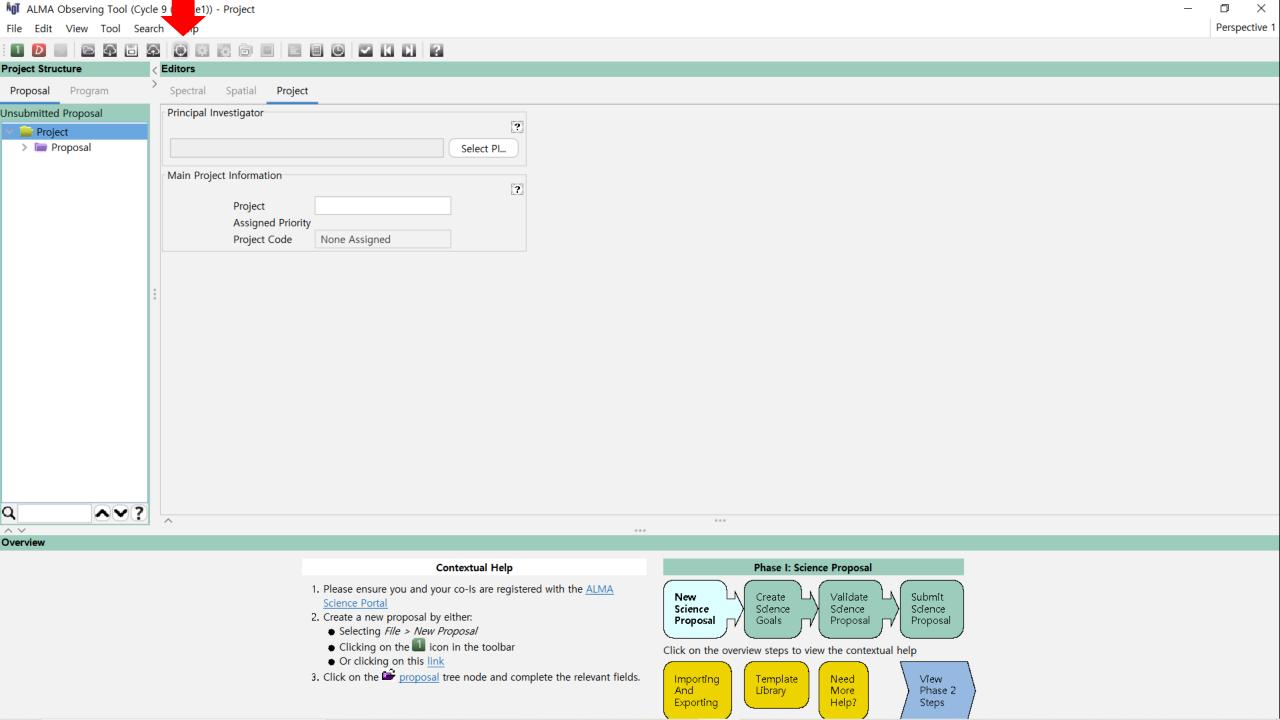
- These values are for each spectral window and for each polarization, using the full correlator resources and no on-line spectral binning.
- Single-polarization modes are also available, giving twice the number of channels per spectral window, and half the channel spacing of the above table.
- 3. The "Channel spacing" is the frequency separation between data points in the output spectrum. The spectral resolution i.e., the FWHM of the spectral response function is larger by a factor that depends on the "window function" applied to the data to control the ringing in the spectrum. For the default function the "Hanning" window this factor is 2. See Chapter 5 of the Technical Handbook for details.
- Only for the 64-input Correlator

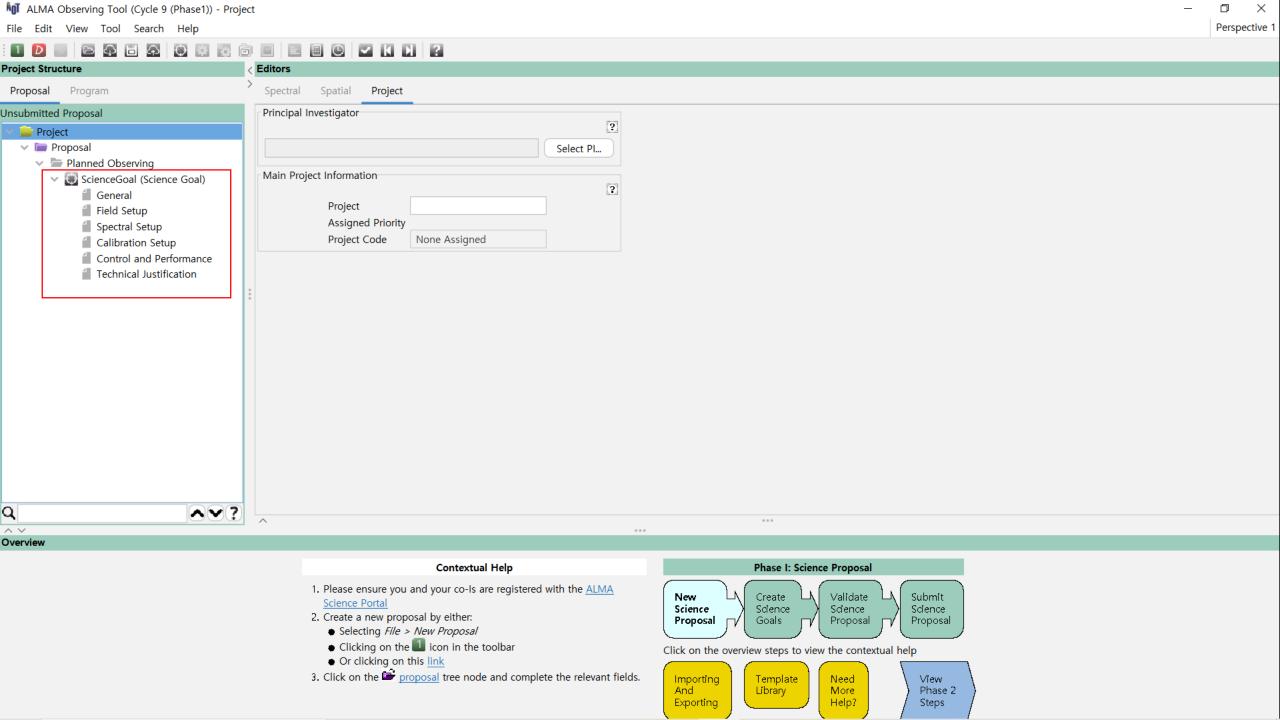
Table 5.1: Available spectral windows in multi-region mode (dual polarization). Each time the fraction is changed, the number of channels and bandwidth of a particular correlator mode is halved. Each row corresponds to a particular spectral resolution.

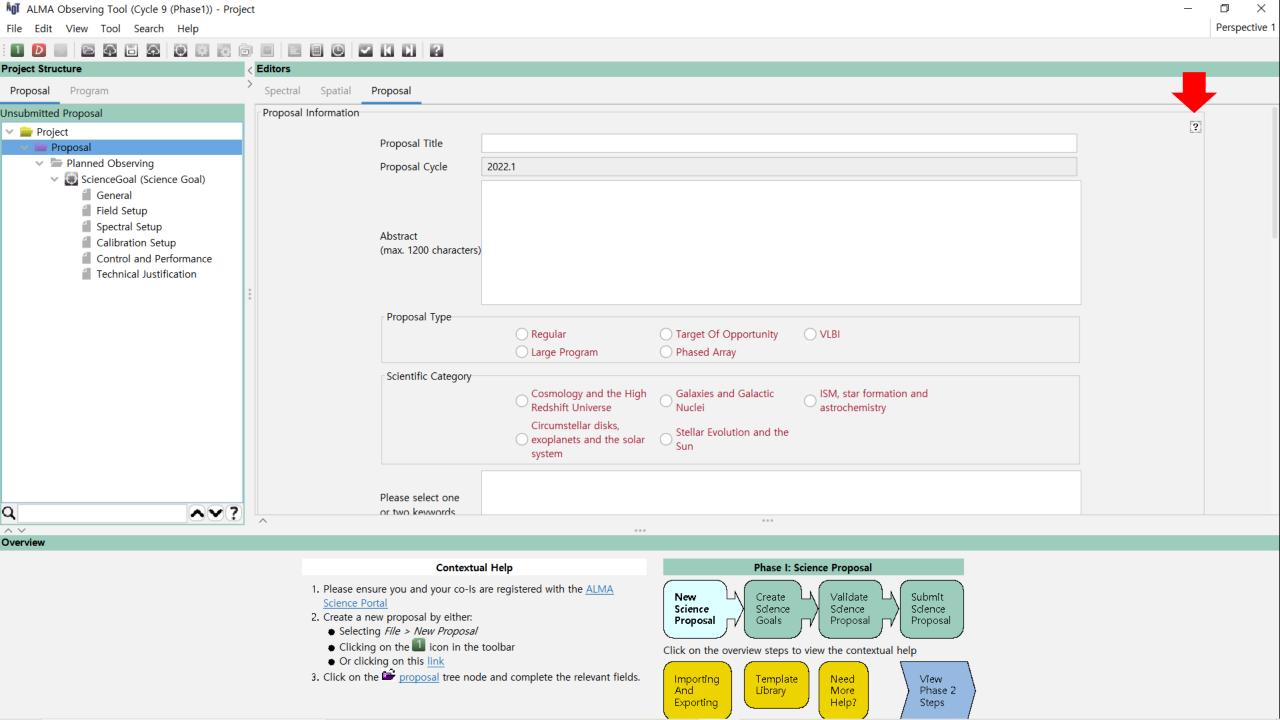
Spectral resolution  $\propto$  1/ fraction for a given bandwidth

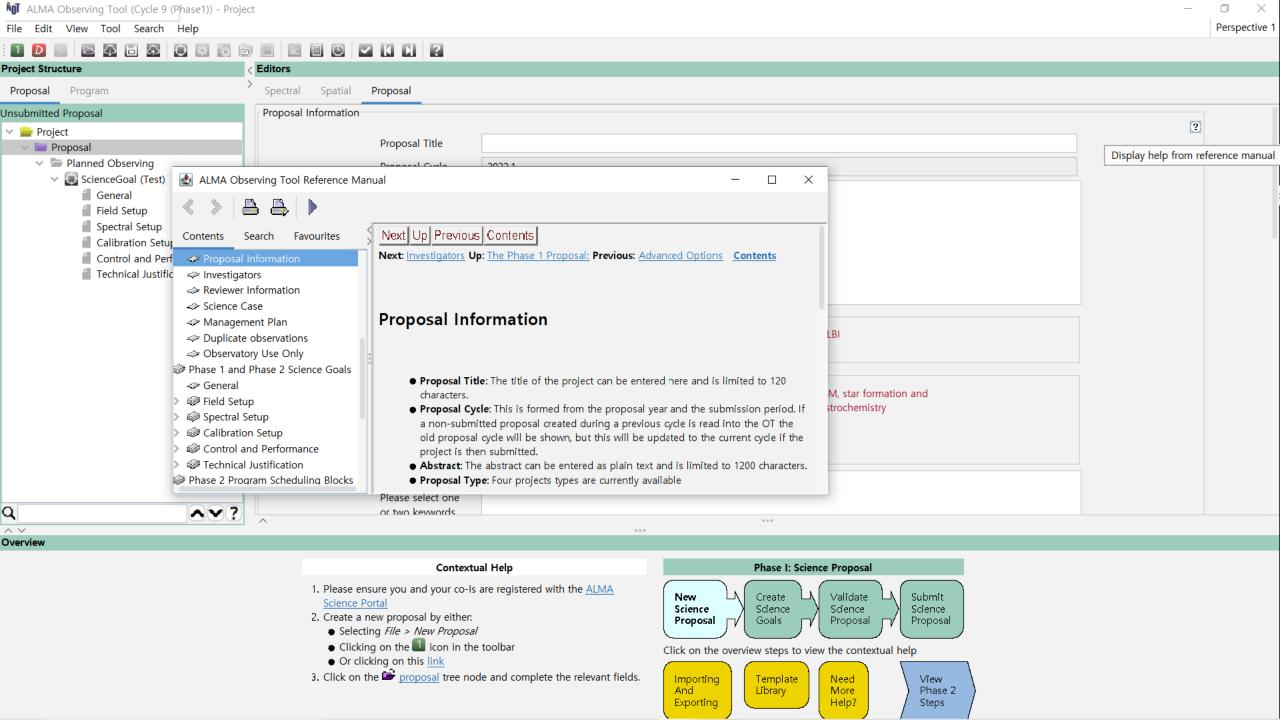
Fraction = 1		Fraction =	1/2	Fraction = 1/4		
Bandwidth (MHz)	# channels	Bandwidth (MHz)	# channels	Bandwidth (MHz)	# channels	
1875	4096	937.5	2048	468.75	1024	
937.5	4096	468.75	2048	234.375	1024	
468.75	4096	234.375	2048	117.118	1024	
234.375	4096	117.118	2048	58.594	1024	
117.118	4096	58.594 2048		not availa	ble	
58.594	4096	not available		not availa	ble	

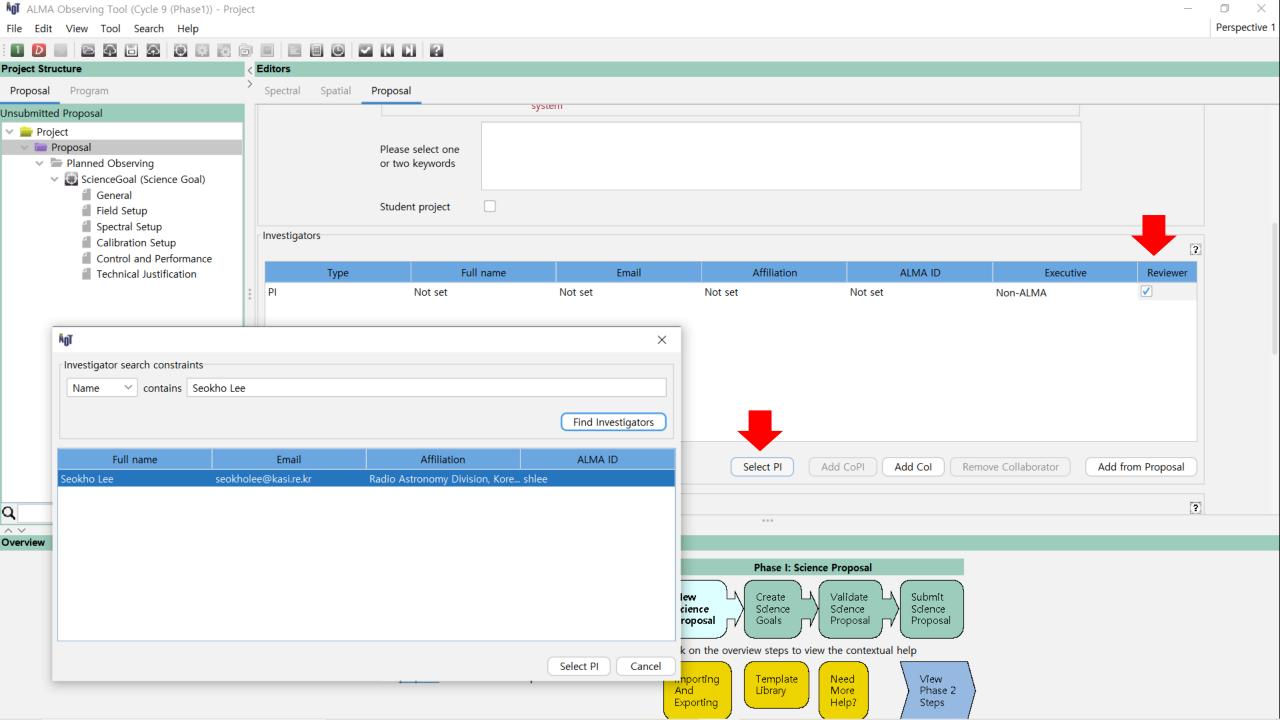
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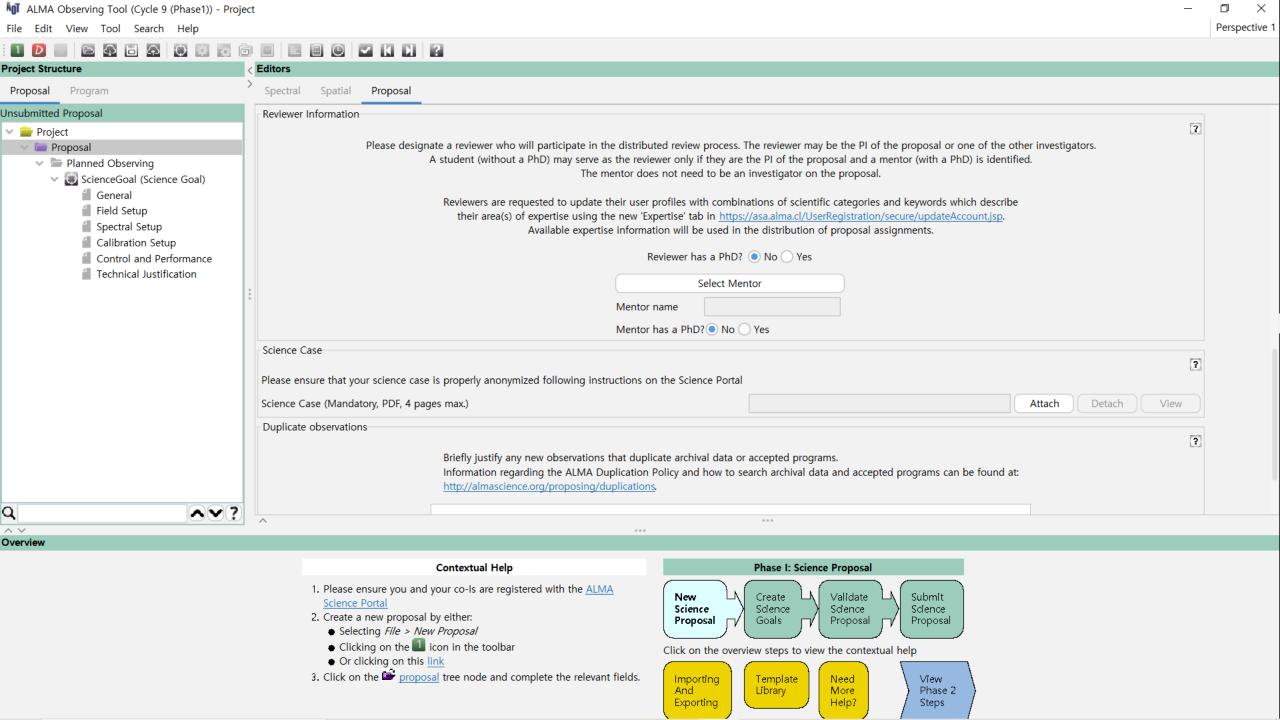


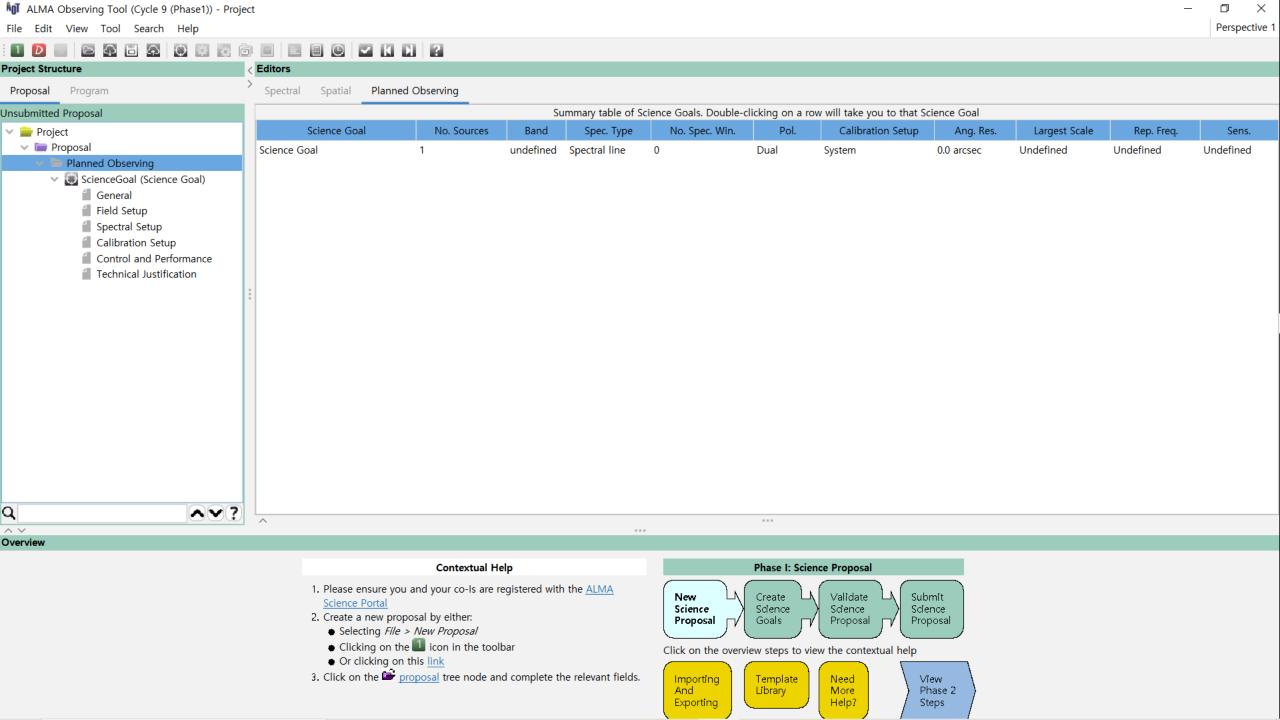


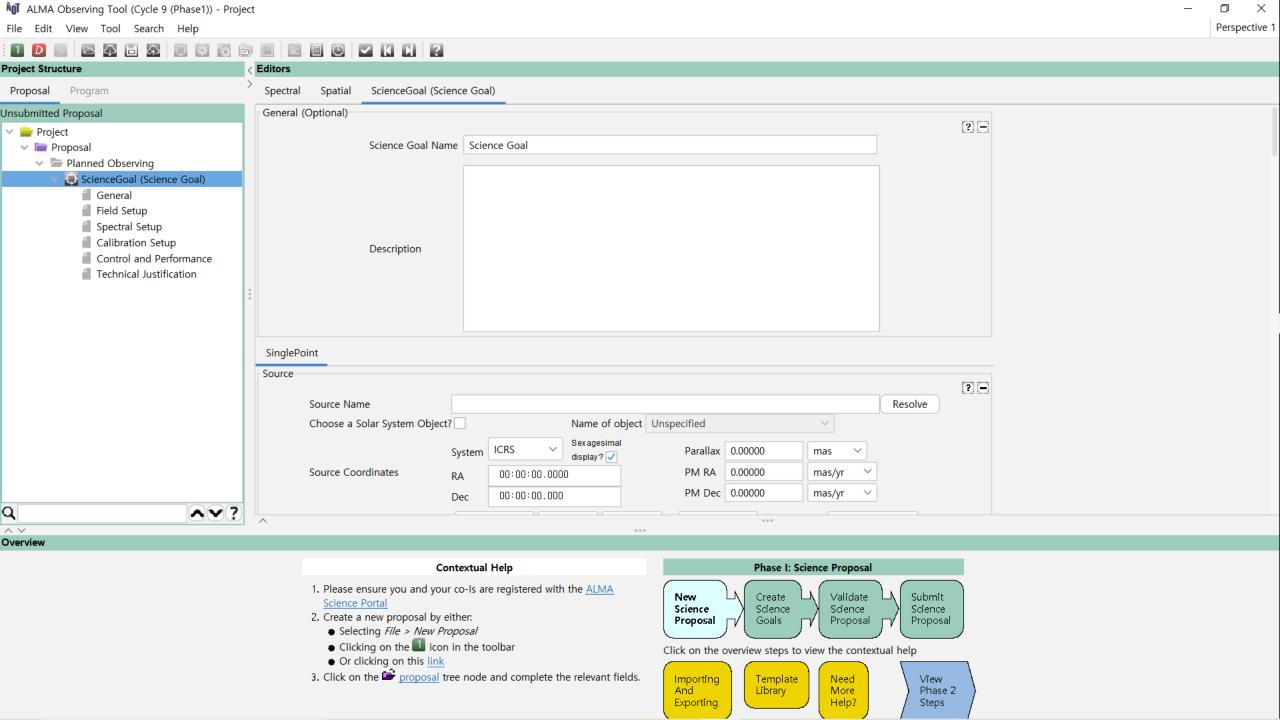


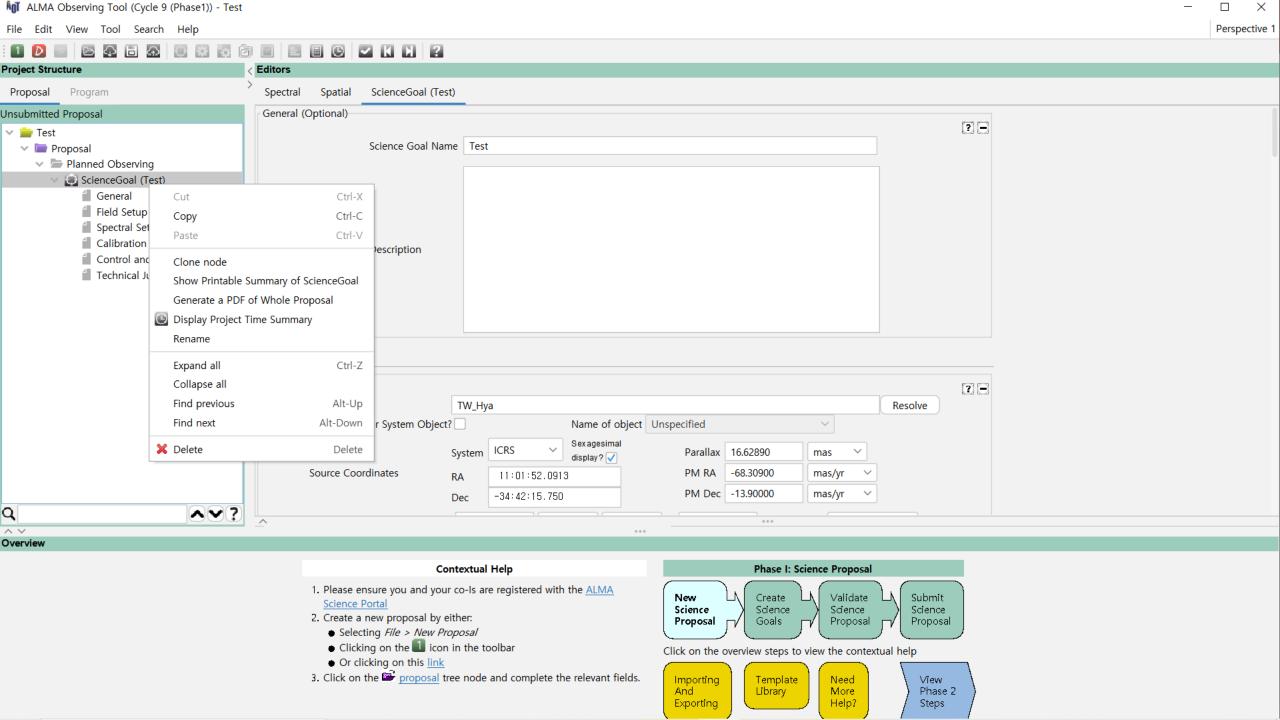


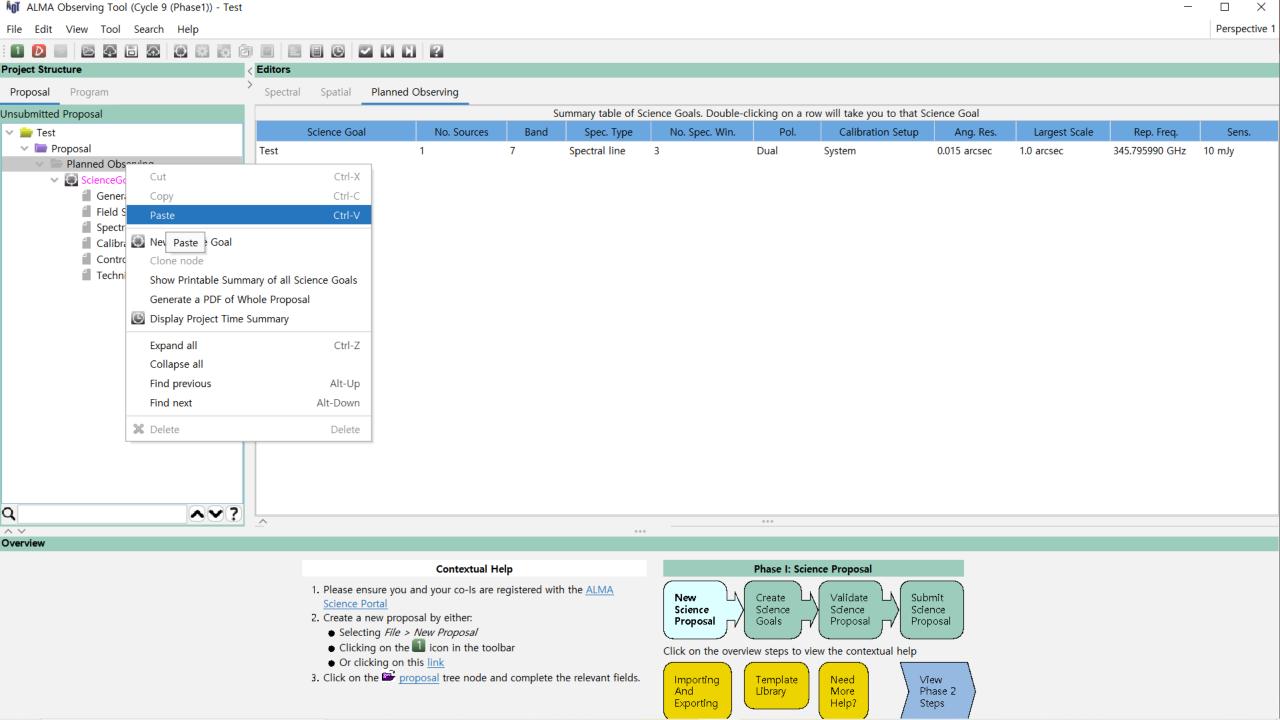


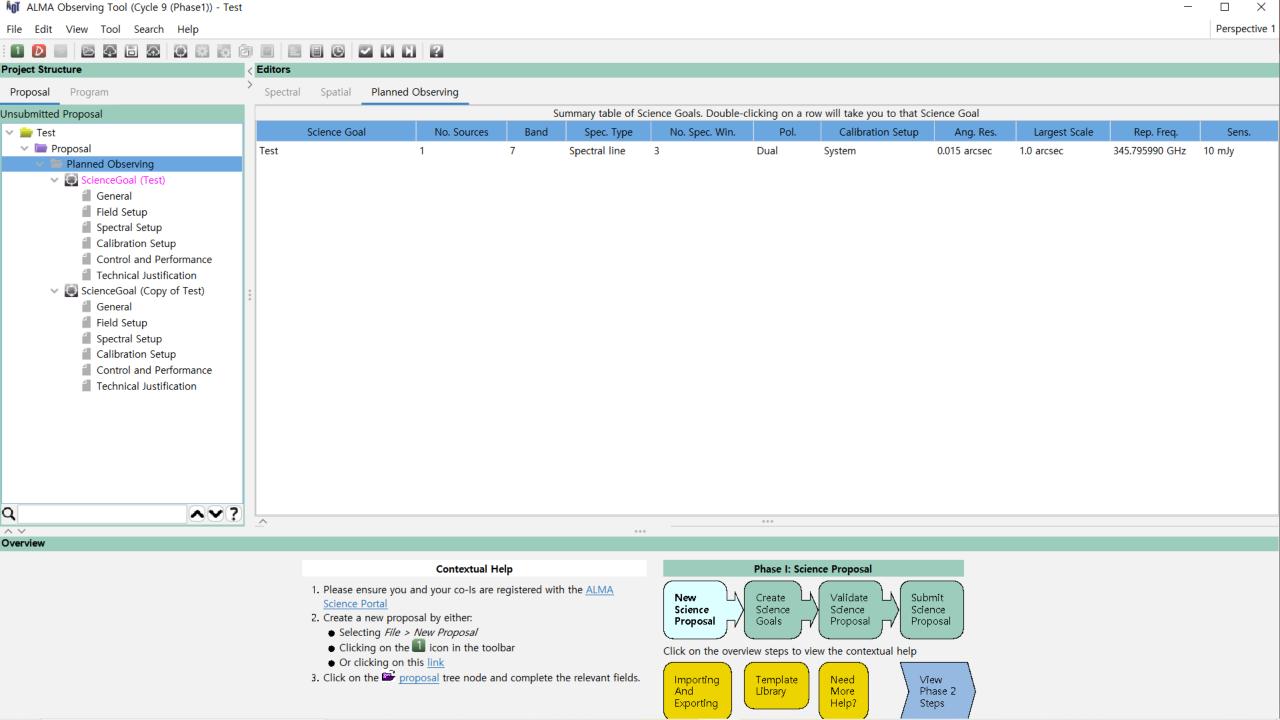


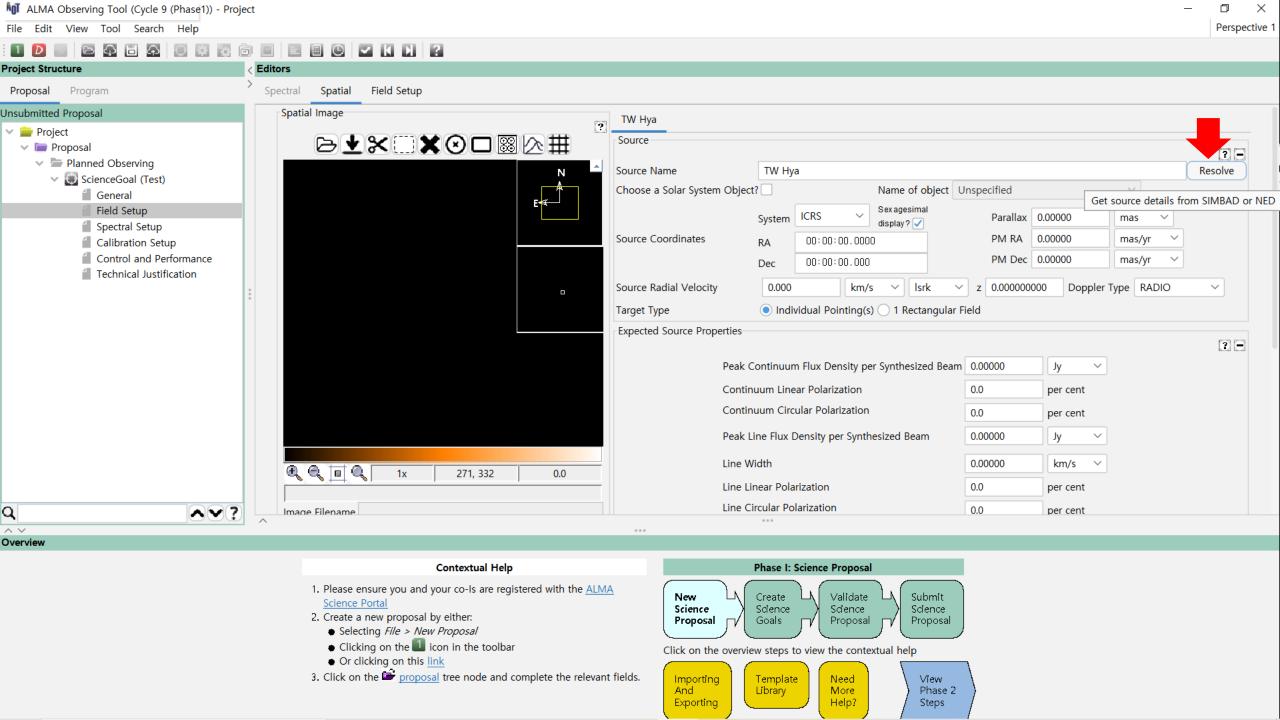


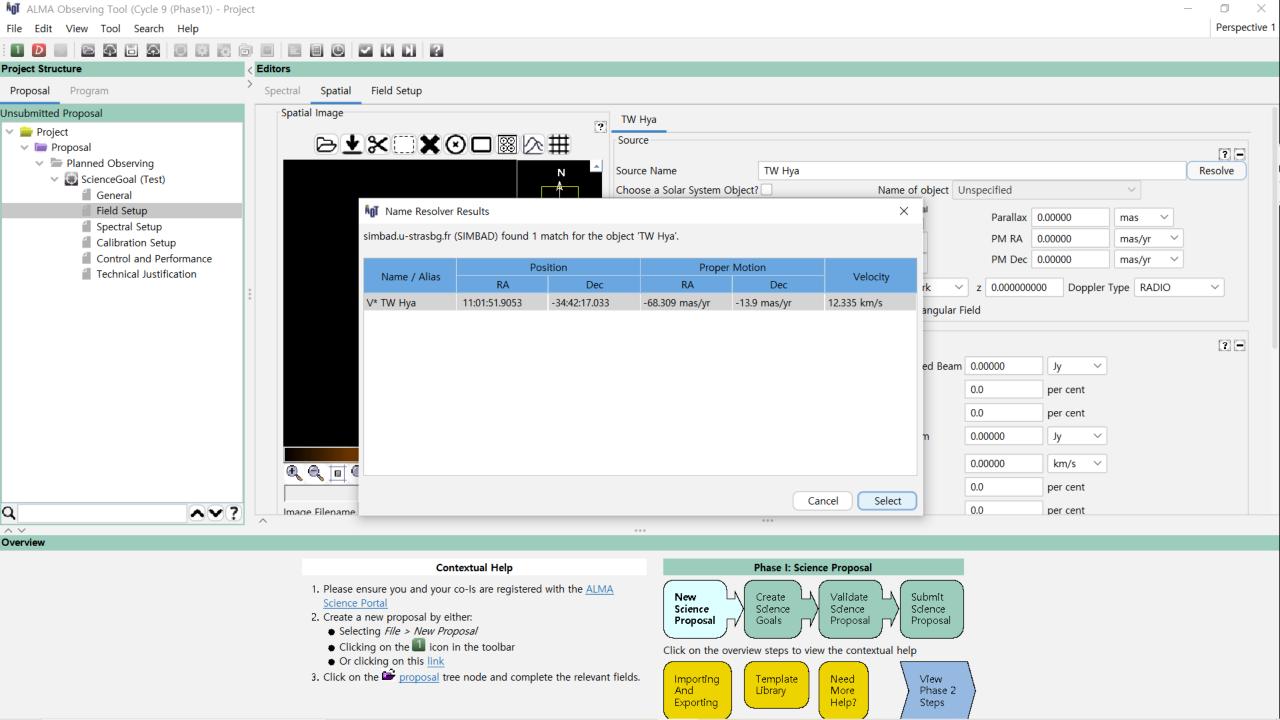


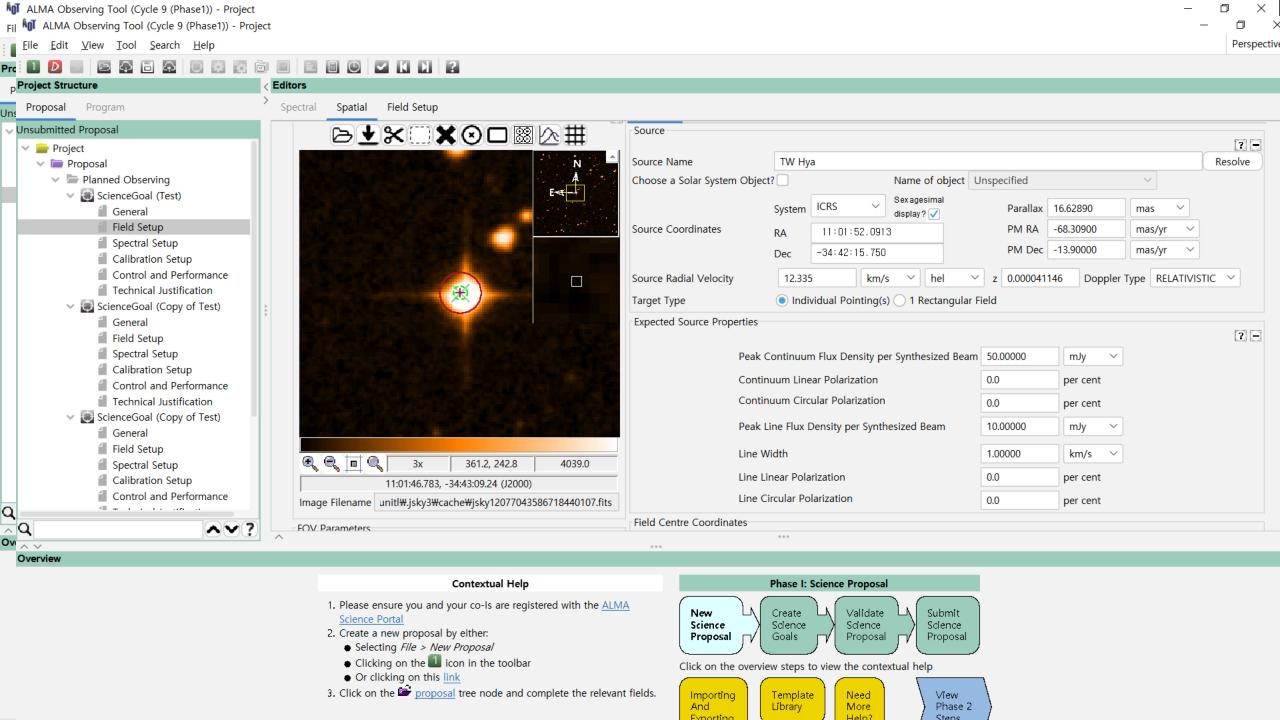


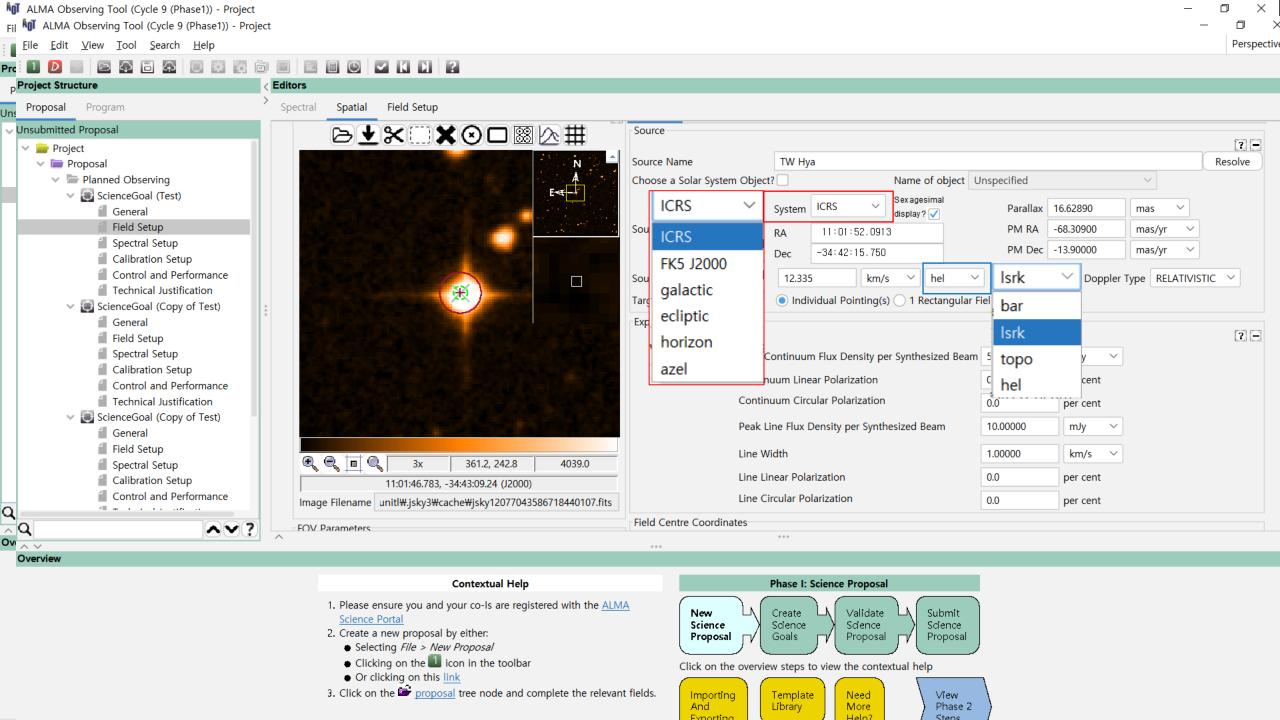


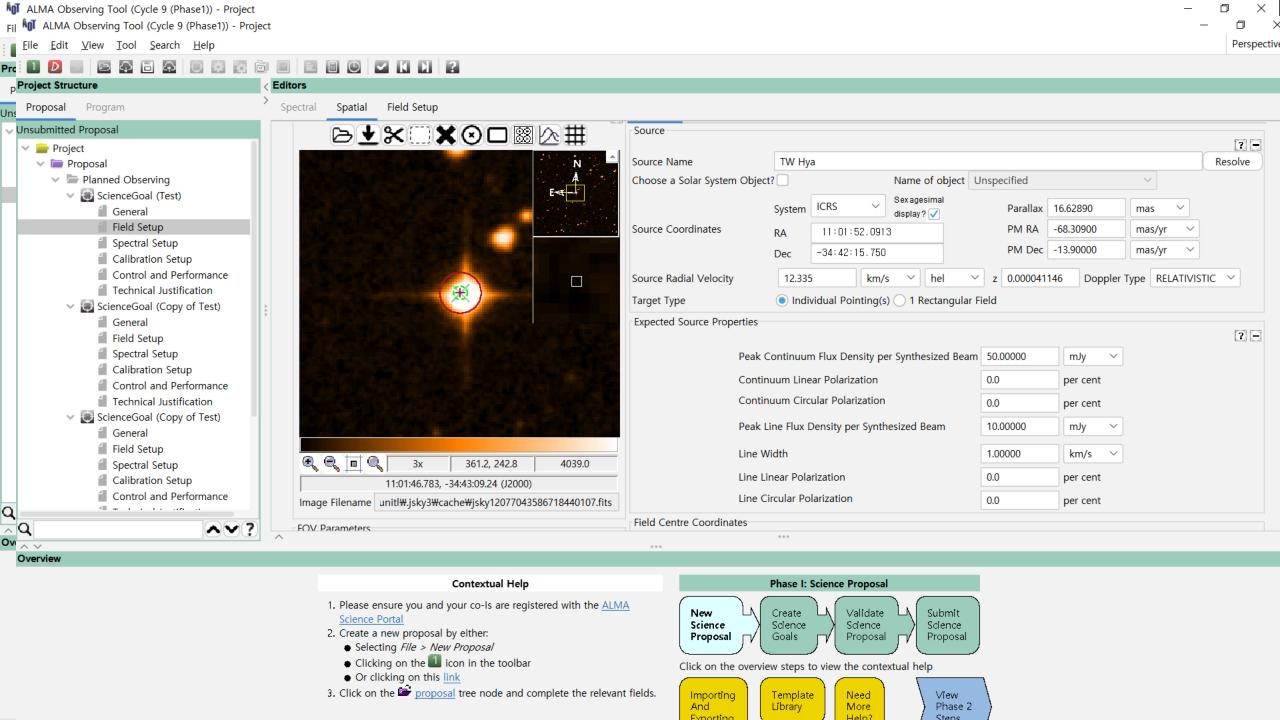


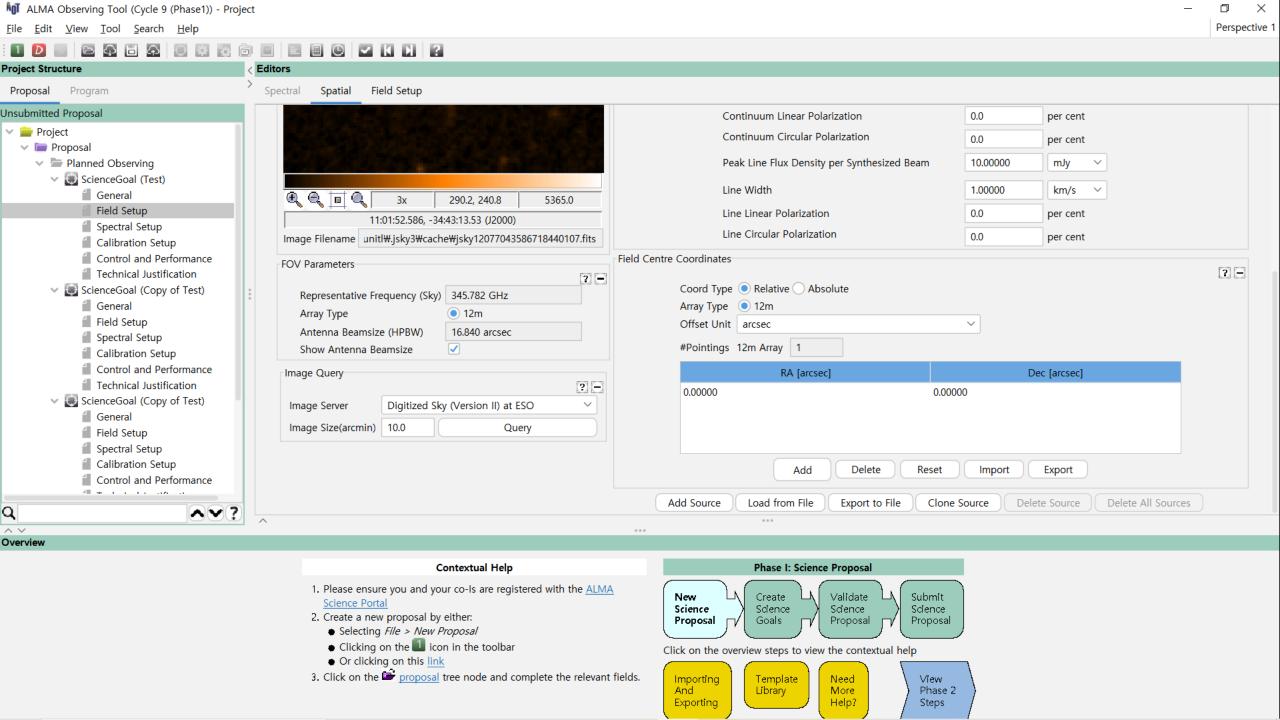


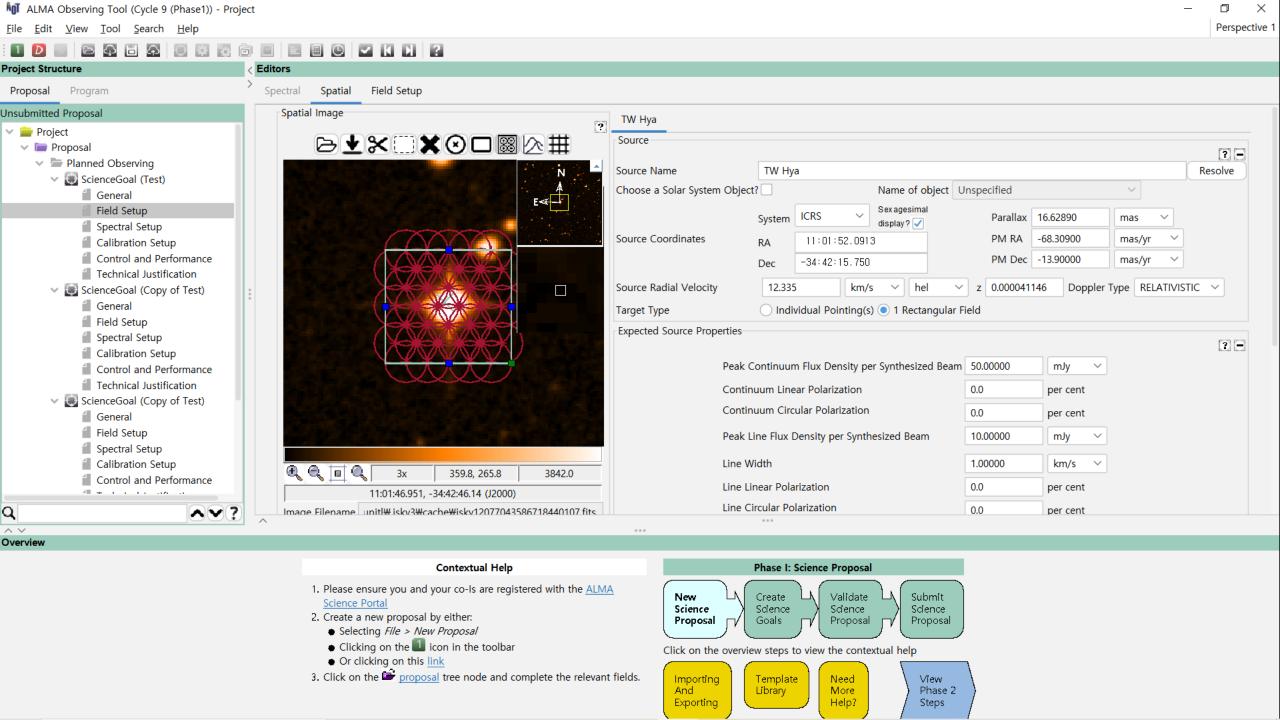


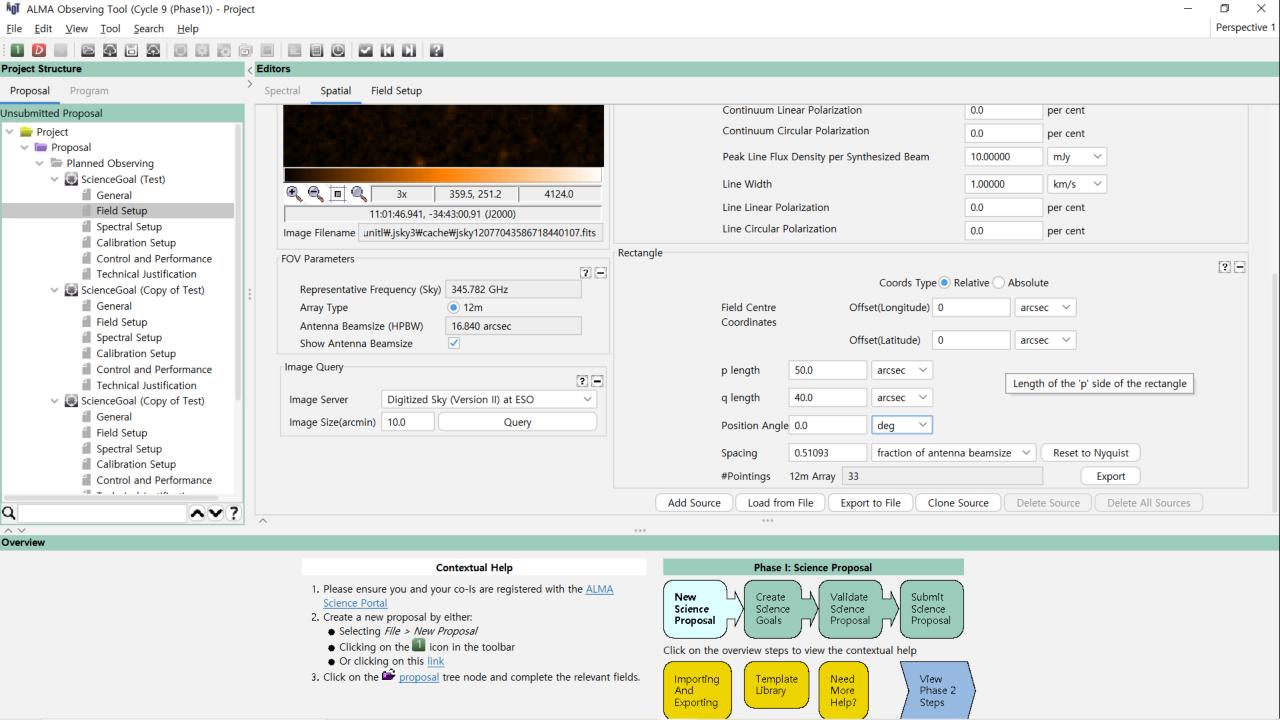


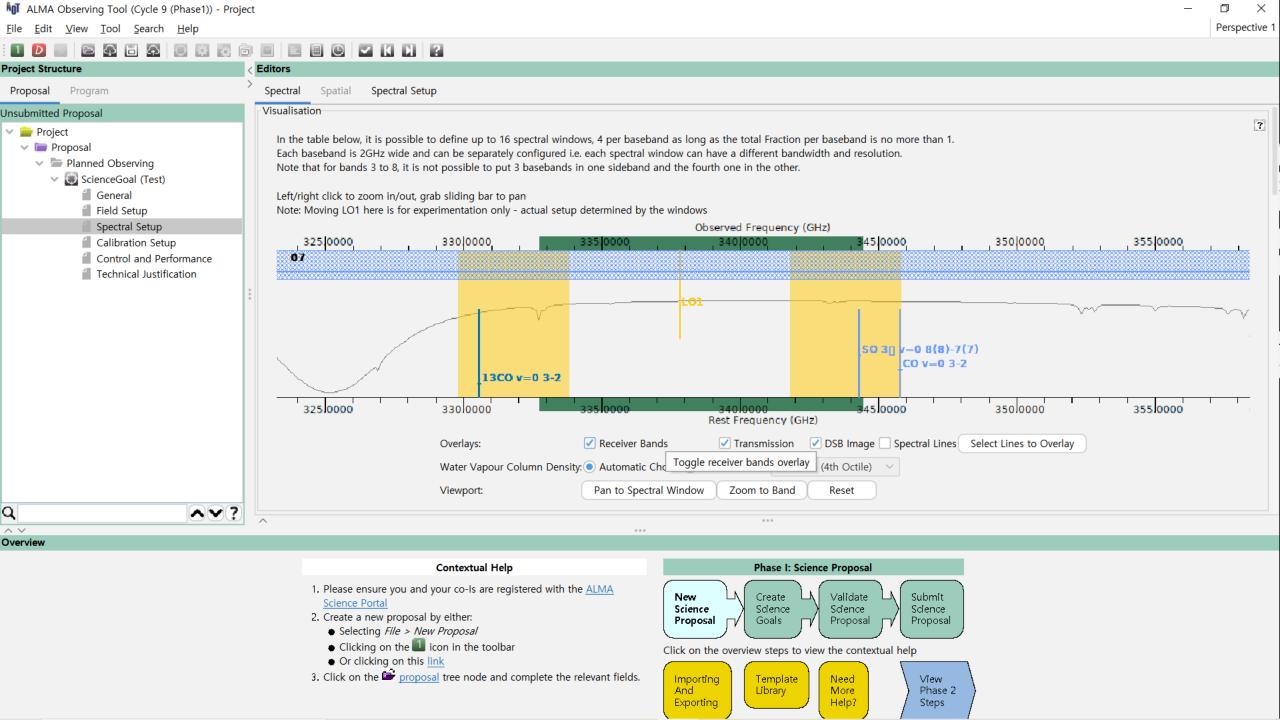


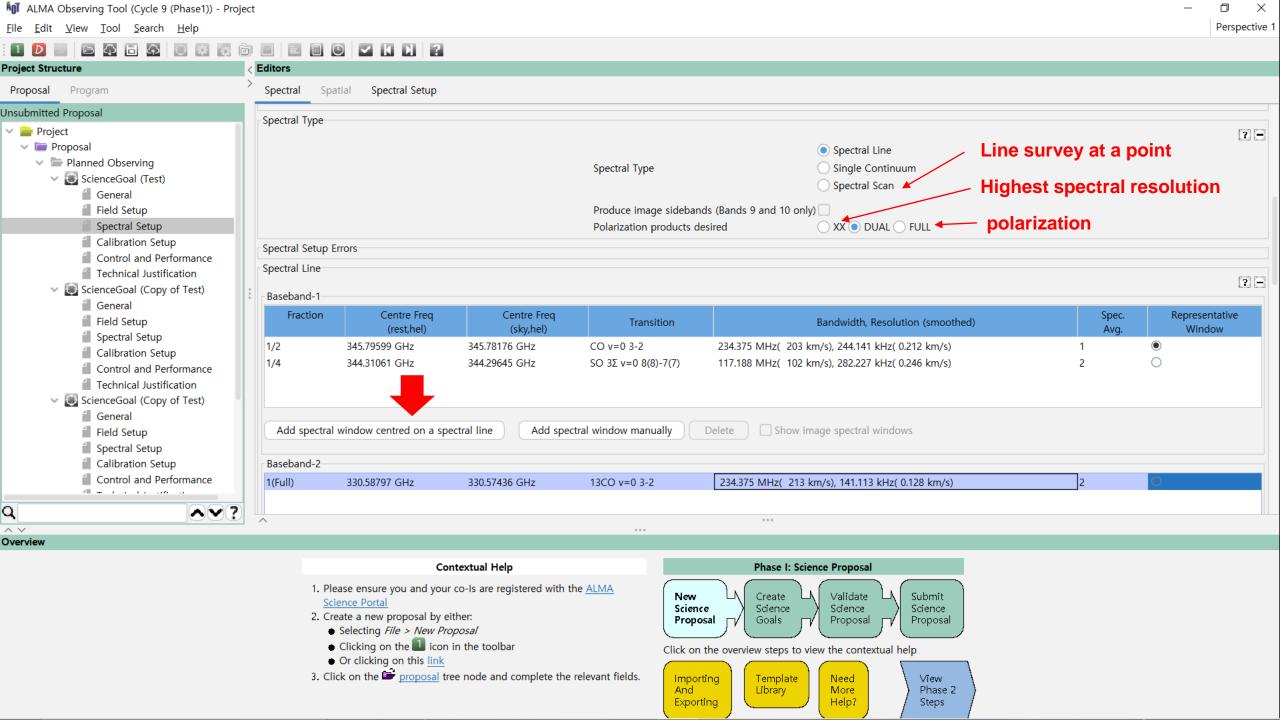


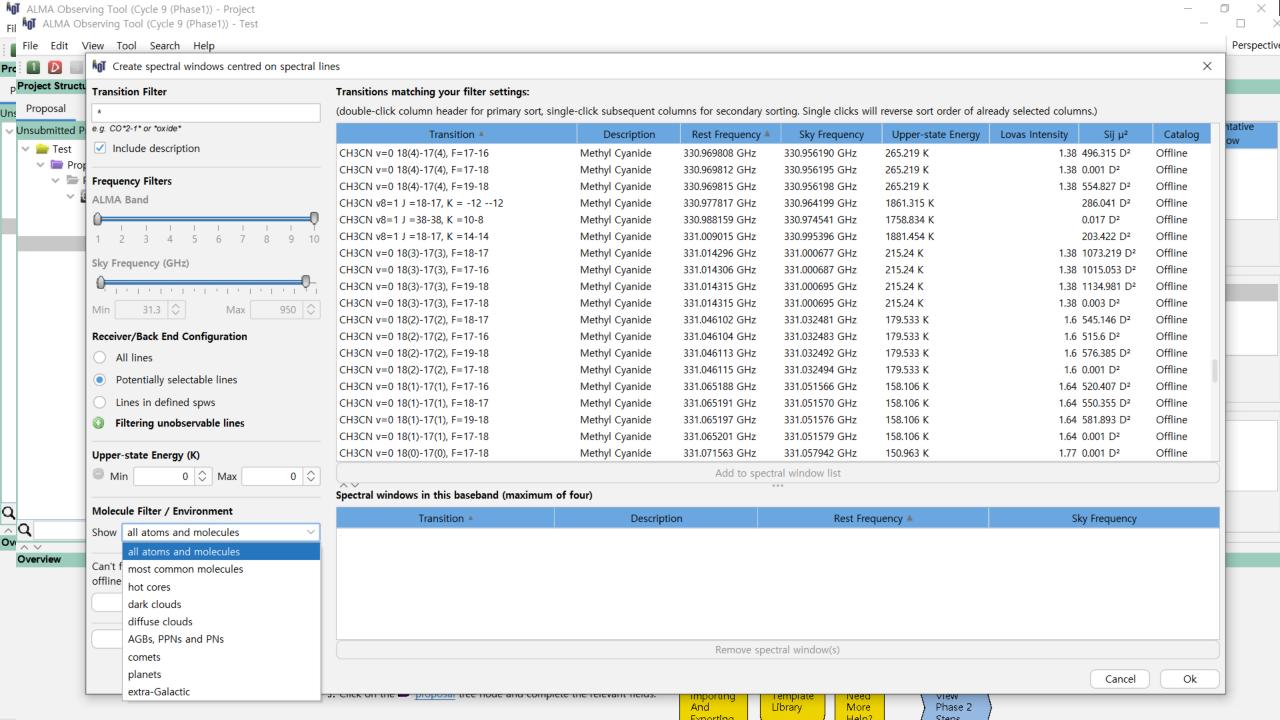


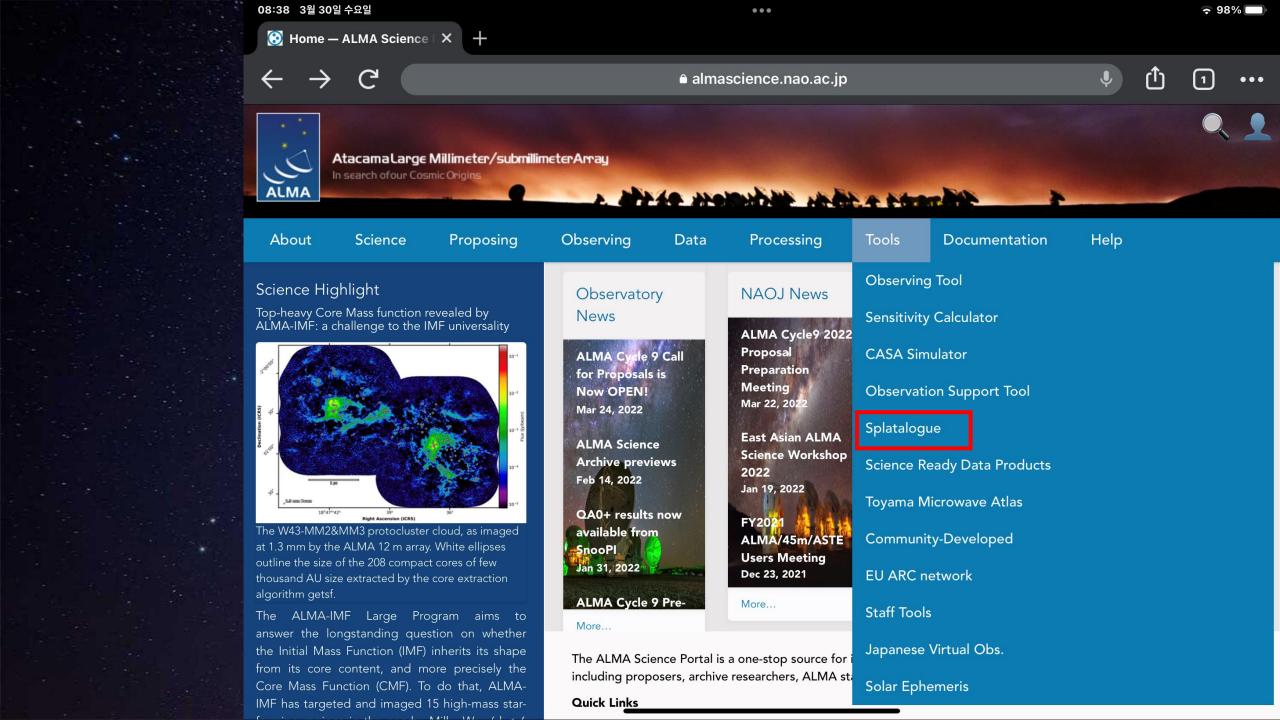


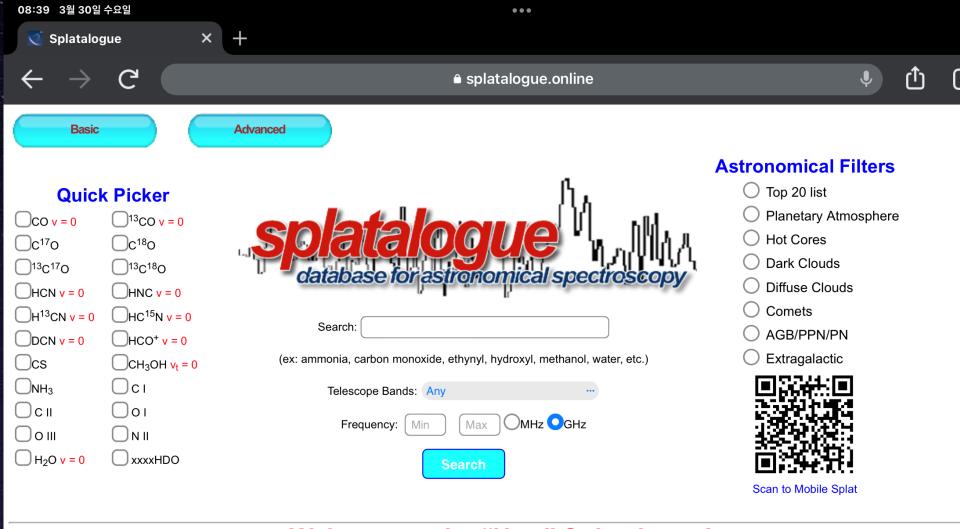












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## Welcome to the "New" Splatalogue!

Over the past several years, there has been an active effort to improve the overall functionality and usability of Splatalogue. We are now offering new options to navigate the nearly 6 million spectral lines available via Splatalogue. The user community has suggested a simpler, more efficient way of searching for and obtaining the more common spectral line features from the radio to submillimeter wavelength.

This new **Splatalogue Basic** search page is now available and has several new and quick search features including:

**The Quick Picker:** Located on the far left. Popular species are included. Click on your favorite, hit search and the results will pop up. You can also limit the frequency by entering in your preferred range in GHz or MHz.

Search Bar: Located in the center of the page. Type in the name (or in some cases, the formula) of your favorite molecule and all species with that

